

# The asynchronous polar V1432 Aquilae and its path back to synchronism

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## **Abstract**

V1432 Aquilae is the only known eclipsing asynchronous polar. In this respect it is unique and therefore merits our attention. We report the results of a 15-year campaign by the globally distributed Center for Backyard Astrophysics to observe V1432 Aql and investigate its return to synchronism. Originally knocked out of synchrony by a nova explosion before observing records began, the magnetic white dwarf in V1432 Aql is currently rotating slower than the orbital period but is gradually catching up. The fortuitously high inclination of the binary orbit affords us the bonus of eclipses providing a regular clock against which these temporal changes can be assessed. At the present rate, synchronism should be achieved around 2100. The continually changing trajectory of the accretion stream as it follows the magnetic field lines of the rotating white dwarf produces a complex pattern of light emission which we have measured and documented, providing comprehensive observational evidence against which physical models of the system can be tested.

## **1. Introduction**

V1432 Aquilae, located at RA 19h 40m 11.42s, Dec  $-10^{\circ} 25' 25.8''$  (J2000), is one of only four known asynchronous polars, the other three being V1500 Cyg, BY Cam and CD Ind. Polars are AM Her type cataclysmic variables (CVs) in which the white dwarf (WD) has a sufficiently strong magnetic field, typically  $>10\text{MG}$ , that formation of an accretion disc is inhibited. Matter transferring from the main sequence star forms an accretion stream which is diverted by the magnetic field of the WD towards one or both of its magnetic poles. Acceleration of the accretion stream onto the surface of the WD releases energy across the electromagnetic

spectrum from X-rays to infrared. In most polars the interaction between the magnetic fields of the WD and the secondary results in rotation of the WD being synchronised with the binary orbital period (Cropper 1990). The common wisdom is that the departure from synchronism seen in these four systems is the result of a relatively recent nova explosion which disrupted the magnetic interaction and that these systems are in the process of slowly returning to synchronism. V1500 Cyg did indeed experience such a nova explosion in 1975. Of the four known asynchronous polars, V1432 Aql is the only one which displays orbital eclipses. These provide a regular clock against which temporal changes in the system can be assessed.

Early observations of this variable optical and X-ray source were attributed to the Seyfert galaxy NGC6814 and caused considerable excitement at the time. It was later identified as a separate source (Madejski et al. 1993) and labelled RX J1940.1-1025 before eventually being given the GCVS designation V1432 Aql.

The object was identified as a likely AM Her type CV by Staubert et al. (1994). The presence of two distinct signals in the optical light curve at  $12116.6 \pm 0.5$ s and  $12149.6 \pm 0.5$ s was first reported by Friedrich et al. (1994). Patterson et al. (1995) confirmed these two periods and noted that the shorter period is the binary orbital period since it is marked by deep eclipses and is consistent with radial velocity measurements. Patterson et al. also suggested that the longer period, which was consistent with a period previously observed in X-rays produced in the accretion process, was the rotation period of the WD thereby identifying V1432 Aql as an asynchronous polar. It was alternatively proposed by Watson et al. (1995) that the eclipse features were occultations of the WD by the accretion stream but subsequent analyses eventually rejected this conclusion. Further photometric and spectroscopic observations at optical and X-ray wavelengths were reported by, among others, Watson et al. (1995), Staubert et al. (2003), Mukai et al. (2003), Andronov et al. (2006) and Bonnardeau (2012).

From spectroscopic data, Watson et al. (1995) concluded that the secondary is a main sequence star of spectral type M4. Friedrich et al. (2000) Doppler-mapped the system and, in the absence of evidence of a distinct accretion stream, concluded that the accretion process is in the form of a curtain over a wide range in azimuth. Mukai et al. (2003) suggested from simple modelling that accretion takes place to both poles at all times and that accreting material may almost surround the WD.

The present analysis includes substantially more data than were available to previous analyses. In section 2 we review our new data. In sections 3, 4 and 5 we analyse the orbital, WD rotation and WD spin periods. In section 6 we review the progress towards resynchronisation and in section 7 we show how observable parameters of the system vary with rotation of the WD. Section 8 identifies some unusual behaviour yet to be explained and section 9 summarises our results.

## 2. Observations

A total of 75849 photometric observations of V1432 Aql were submitted to the Center for Backyard Astrophysics (CBA) in 312 datasets by 23 observers between 1998 and 2013. These comprised a total of 1170 hours of observation. A breakdown by year is given in Table 1.

Year	Start date	End date	No of runs	Total time (hrs)
1998	06 May	23 Sep	13	26
1999	18 May	06 Jun	6	12
2000	21 May	27 Oct	12	30
2002	26 Jul	27 Sep	75	346
2007	07 Jul	20 Sep	52	221
2008	28 Jun	22 Jul	2	6
2011	26 Jun	01 Nov	45	165
2012	07 Jun	29 Sep	68	244
2013	03 Jun	11 Nov	39	120

**Table 1. Summary of observations.**

The CBA (<http://cbastro.org/>) is a globally distributed network of small telescopes operated by amateur astronomers interested in cooperating in the study of variable stars. Almost all observations were obtained unfiltered to maximise time resolution and signal-to-noise and used a variety of CCD cameras and telescopes in the aperture range 0.2-m to 0.4-m. All images were bias- and dark-subtracted and flat-fielded before differential magnitudes were measured with respect to nearby comparison stars. Observers used a range of comparison stars, although most observers remained loyal to one comparison star, so alignment in magnitude between datasets had to be achieved empirically. By an iterative process, datasets from different observers which overlapped in time were aligned in magnitude and then others close in time were aligned with those. Eventually all datasets were brought into alignment in magnitude with an uncertainty estimated to be less than 0.1 magnitude, thus generating an integrated and internally consistent light curve spanning 15 years for use in our subsequent analysis. All observation times were converted to Heliocentric Julian Dates (HJD).

## 3. Orbital period

The most recent published orbital ephemeris from Bonnardeau (2012) was used to assign a provisional orbital phase to every observation. Segments of light curves between orbital phases  $-0.05$  and  $+0.05$  were extracted and a second order polynomial fitted to the minimum of each eclipse. Eclipses were generally symmetrical and round-

bottomed and were well-fitted by a quadratic from which we could find the time and magnitude at minimum. In cases where the eclipses were significantly asymmetrical, only those points close to the minimum were used to ensure we found an accurate time of minimum. Eclipses poorly defined due to a large scatter in magnitude were rejected. An estimate was made of the average uncertainty in the time of minimum for each observer during each year by computing the root mean square residual between their times of minimum and a linear ephemeris determined for all measured eclipses in that year. This estimated uncertainty was then assigned to all observations by that observer in that year. The process was iterated to test its stability. Orbital cycle numbers were assigned to each eclipse based on the mid-eclipse ephemeris in Patterson et al. (1995).

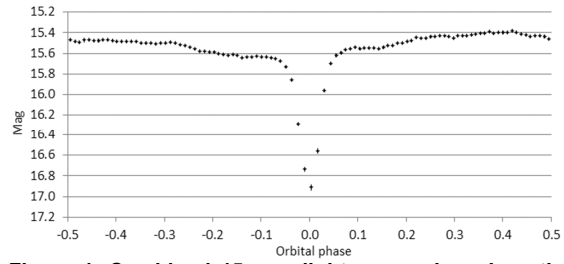
At this stage we cannot say that these times represent the true times of mid-eclipse of the WD by the secondary star since there may be sources of light external to the WD which also experience eclipse and could alter the shape of the light curve around the true time of mid-eclipse. However, for the following analysis of the orbital period, we assume that these other effects average out over time and we take our derived times of minimum as times of mid-eclipse. Orbital cycle numbers, mid-eclipse times and assigned uncertainties for 228 measured eclipses are listed in Table 2.

A further 121 eclipse times including optical and X-ray observations were extracted from published papers by Patterson et al. (1995), Watson et al. (1995), Mukai et al. (2003), Andronov et al. (2006) and Bonnardeau (2012). In some cases it was clear that the published uncertainties on these times bore little relation to their intrinsic scatter and therefore could not be used to compute reliable weights on these times. For this reason, uncertainties were recomputed as described above for the times published in each paper.

We computed the following orbital ephemeris as a function of the orbital cycle number  $E$  by a weighted linear regression using all 349 times of mid-eclipse. Each mid-eclipse time was weighted by the inverse square of its assigned uncertainty.

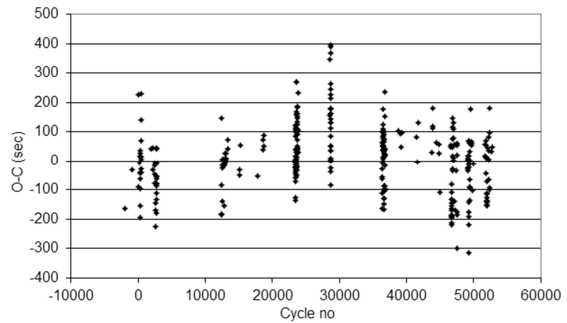
$$\text{HJD (mid-eclipse)} = 2449199.69307(6) + 0.140234751(3).E \quad (1)$$

We adopt the orbital period  $P_{\text{orb}} = 0.140234751\text{d}$  (12116.282s) in our subsequent analysis. Our combined 15-year light curve phased on this orbital period and averaged in 100 bins is shown in Figure 1.



**Figure 1. Combined 15-year light curve phased on the orbital period 0.140234751d and averaged in 100 bins.**

The O-C (Observed minus Calculated) residuals to this linear ephemeris are shown in Figure 2. The rms scatter is 110s and the reduced chi-squared of this fit is 1.02.

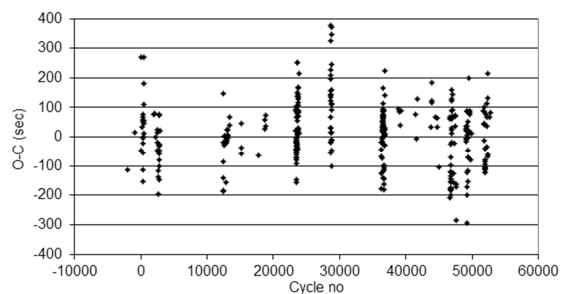


**Figure 2. O-C residuals of eclipse times to the linear orbital ephemeris in equation (1).**

The weighted eclipse times were also fitted with a 2nd order polynomial which gave the following quadratic orbital ephemeris.

$$\text{HJD (mid-eclipse)} = 2449199.69257(12) + 0.140234804(12).E - 9.7(2.1) \times 10^{-13}.E^2 \quad (2)$$

The O-C residuals to this quadratic ephemeris are shown in Figure 3. This ephemeris represents a rate of change of orbital period  $dP_{\text{orb}}/dt = -1.38(29) \times 10^{-11}$  years/year. The rms scatter is 105s and the reduced chi-squared is 0.96.



**Figure 3. O-C residuals of eclipse times to the quadratic orbital ephemeris in equation (2).**

#### 4. WD rotation period

Previous analyses have concluded that the WD is slowly rotating in the rest frame of the binary system about an axis perpendicular to the orbital plane with a period which is currently about 62 days. This period is slowly lengthening due to interaction between the magnetic field of the WD and the accretion stream from the secondary which is creating a torque slowing the WD rotation. In time the WD will stop rotating relative to the secondary and the system will have resynchronised. As seen from our vantage point (which orbits around the rest frame of the binary system every 3hr 22min) the WD will then appear to spin in synchrony with the orbital period of the binary.

The WD rotation period in the rest frame of the binary,  $P_{\text{rot}}$ , is given by the beat period between the orbital period  $P_{\text{orb}}$  and apparent WD spin period  $P_{\text{spin}}$  as seen from our orbiting vantage point.

$$1/P_{\text{rot}} = 1/P_{\text{orb}} - 1/P_{\text{spin}}$$

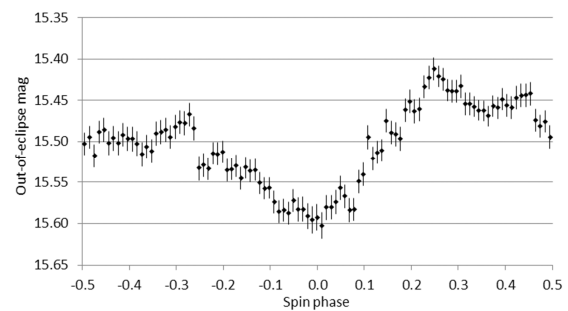
To predict when resynchronisation will occur we need to measure  $P_{\text{spin}}$  and hence find  $P_{\text{rot}}$  and the rate at which it is changing.

#### 5. WD spin period

Geckeler & Staubert (1997) and Staubert et al. (2003) proposed an accretion scenario based on a dipole model of the WD magnetic field. This predicts modulation in the out-of-eclipse light curve due to the changing aspect of the accretion spot on the surface of the rotating WD. To search for that signal, we first removed the eclipses in the light curve between orbital phases  $-0.05$  and  $+0.05$ . We then carried out a period analysis of the remaining out-of-eclipse light curve year by year. This gave initial estimates of the WD spin period for each year. The uncertainty on these estimates was too large to accurately define a rate of change but they were sufficient to phase the out-of-eclipse light curve for each year on this spin period. As noted by several authors, there is a prominent ‘spin dip’ or ‘trough’ in the spin-phased light curve for each year. We took the average position of the spin dip for each year as defining spin phase 0.0 in that year. Staubert et al. (2003) suggested that this dip occurs due to absorption of light by the accretion column. The timing of these dips potentially provides a way of measuring the WD spin period accurately. We note in passing that there is possible confusion in

nomenclature as some papers refer to the eclipse as a ‘dip’.

Figure 4 shows the combined 15-year out-of-eclipse light curve phased within each year on the WD spin period for that year and averaged in 100 bins. We performed a period analysis on the whole 15 year out-of-eclipse dataset. As might be expected given the variation in spin period over this interval, the power spectrum contained many signals of similar strength spanning a broad range of spin period and provided little useful information. No signal was detectable at the orbital period.



**Figure 4. Combined 15-year out-of-eclipse light curve phased within each year on the WD spin period for that year and averaged in 100 bins.**

The times of minimum of all sufficiently well-defined dips within the spin phase range  $-0.1$  to  $+0.1$  were measured and uncertainties assigned as described above for eclipses. A total of 172 spin dip times were measured and are listed in Table 3. A further 73 spin dip times were obtained from the published literature and uncertainties on these reassigned as before. A preliminary linear ephemeris derived from these 245 times was used to assign an initial cycle number to each spin dip. The resulting O-C residuals revealed a strongly quadratic behaviour but also evidence that our initial cycle number assignment was not perfect. The initial cycle number assignment was adjusted until we achieved smoothly-varying and internally consistent O-C residuals. This process was repeated several times with different initial linear ephemerides to check that the set of assigned cycle numbers was stable and, we therefore assume, correct. These spin cycle numbers are also listed in Table 3. The O-C residuals to a linear ephemeris are shown in Figure 5(upper). We then calculated a quadratic ephemeris whose O-C residuals are shown in Figure 5(middle). Clearly this ephemeris, while an approximate match to the data, is still not a particularly good fit. It does however provide an average value for the rate of change of the WD spin period  $dP_{\text{spin}}/dt = -1.042(1) \times 10^{-8}$  years/year. This is consistent with values found by Staubert et al.

(2003) and Mukai et al. (2003) among others. By increasing the order of the ephemeris, we found that the following fifth order polynomial in the spin cycle number  $E$  provided the best fit to the data and that this could not be improved with a higher order fit.

$$\text{HJD (spin dip)} = \alpha + \beta.E + \gamma.E^2 + \delta.E^3 + \epsilon.E^4 + \zeta.E^5 \quad (3)$$

$$\begin{aligned} \text{where } \alpha &= 2448921.5330(23) \\ \beta &= 0.14063204(75) \\ \gamma &= -2.09(76) \times 10^{-10} \\ \delta &= -2.83(33) \times 10^{-14} \\ \epsilon &= 5.65(64) \times 10^{-19} \\ \zeta &= -3.83(45) \times 10^{-24} \end{aligned}$$

The O-C residuals to this fifth order ephemeris are shown in Figure 5(lower). We are not claiming any physical justification for this degree of fit, simply that it provides the closest polynomial fit to the data and therefore the best basis for determining the varying WD spin and rotation periods. The reduced chi-squared for this fit was 1.00, more than a factor of two better than any lower order fit.

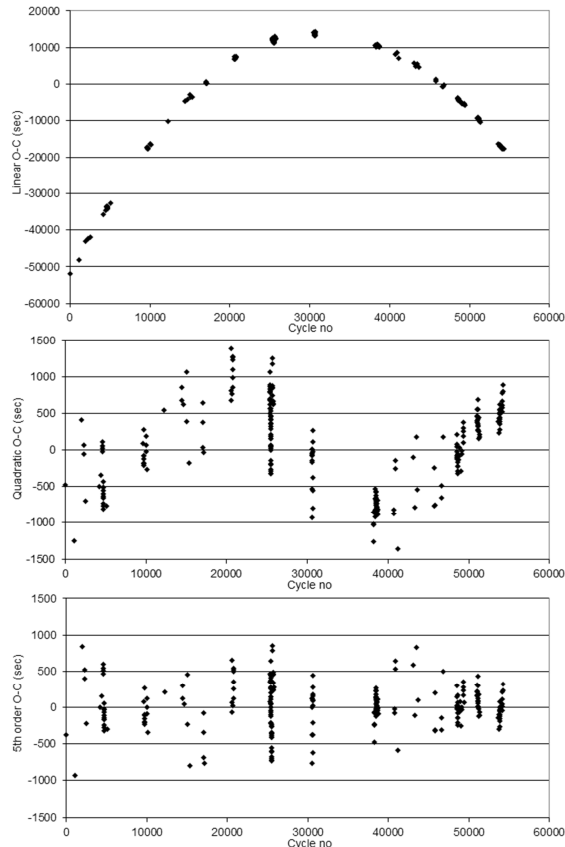


Figure 5. O-C residuals of spin dip times to a linear ephemeris (upper), to a quadratic ephemeris (middle), and to the best fit ephemeris in equation (3) (lower).

We also investigated an exponential fit of the form  $a.\exp(b.x)$  to the spin dip times. This provided a poor fit, quite close to the initial linear ephemeris above.

Equation (3) gives the times of spin dips as a polynomial in the spin cycle number  $E$ . By differentiating this polynomial with respect to  $E$ , we obtained the following expression for the spin period as a function of  $E$ .

$$P_{\text{spin}} = \beta + 2.\gamma.E + 3.\delta.E^2 + 4.\epsilon.E^3 + 5.\zeta.E^4 \quad (4)$$

We calculated the value of  $P_{\text{spin}}$  at each spin dip time and fitted these with a fifth order polynomial in HJD. This provided a means to compute  $P_{\text{spin}}$  and hence  $P_{\text{rot}}$  at any time. We chose  $\text{HJD} = 2449544.81070$ , the mid-eclipse time closest to a spin dip minimum, as defining WD rotation phase 0.0. By numerical integration in 1 day steps, we determined the WD rotation phase at the time of every observation, eclipse and spin dip. The integration step size was determined empirically by reducing the step size until the computed WD rotation phase at the end of 15 years remained stable.

## 6. Progress towards resynchronization

We found that  $P_{\text{spin}}$  reduced from 0.1406320d (12150.61s) in October 1992 to 0.1405542d (12143.88s) in September 2013. The corresponding increase in  $P_{\text{rot}}$  was from 49.641d to 61.703d. Figure 6 illustrates these changes.

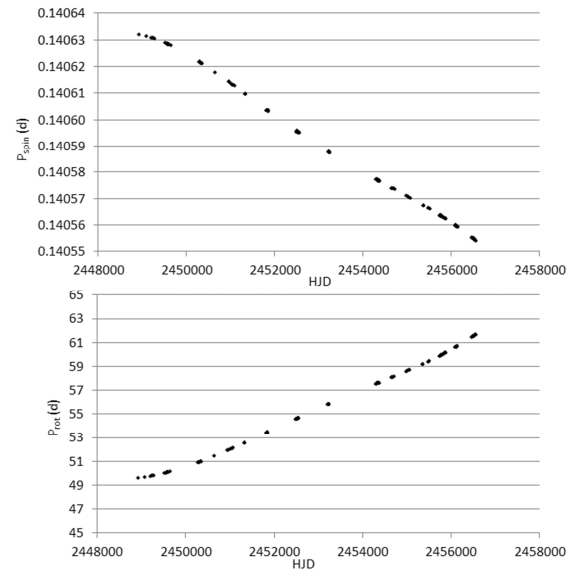


Figure 6. Time variation of  $P_{\text{spin}}$  (upper) and  $P_{\text{rot}}$  (lower).

During this interval the WD rotated 139 times in the binary rest frame. These results are consistent with values of  $P_{\text{spin}}$  found at various times by Patterson et al. (1995), Geckeler & Staubert (1997), Staubert et al. (2003) and Andronov et al. (2006).

In the other three asynchronous polars the apparent WD spin period  $P_{\text{spin}}$  is shorter than the orbital period  $P_{\text{orb}}$ . This is generally assumed to be because the WD has been spun up during a recent nova explosion and is subsequently slowing back into synchrony. Why the WD spins slower in V1432 Aql is an unanswered question.

We can extrapolate the trend in  $P_{\text{rot}}$  between 1992 and 2013 to find when resynchronisation is likely to occur. This will happen when  $1/P_{\text{rot}}$  becomes zero. Using all the currently available data, linear extrapolation predicts resynchronisation will occur in 2097 while quadratic extrapolation predicts it will occur in 2106. This is of the same order as the resynchronisation timescale of  $\sim 170$  years found for V1500 Cyg (Schmidt et al. 1995).

## 7. Variation with WD rotation phase

We investigated how various observable properties of the system vary with the WD rotation phase in the rest frame of the binary. This provides observational evidence against which physical models of the system may be tested.

### 7.1 Variation of orbital light curve with WD rotation phase

In Figure 7 we show the averaged orbital-phased light curve in each of 10 bins of WD rotation phase. The data plotted in this and subsequent similar plots are the mean values for each bin and the error bars represent the standard error for each bin. The light curves show a broad minimum which moves forward in orbital phase as the WD rotation phase increases. Plotting the orbital phase of this minimum against the corresponding WD rotation phase (Figure 8) shows a clear linear relationship. To investigate this further we removed the eclipses and plotted the averaged out-of-eclipse light curve against the difference between the orbital and WD rotation phases (Figure 9). This highlights the effect and shows there is a drop of 0.7 magnitudes, corresponding to a flux drop of about 50%, when the orbital and the WD rotation phases are the same.

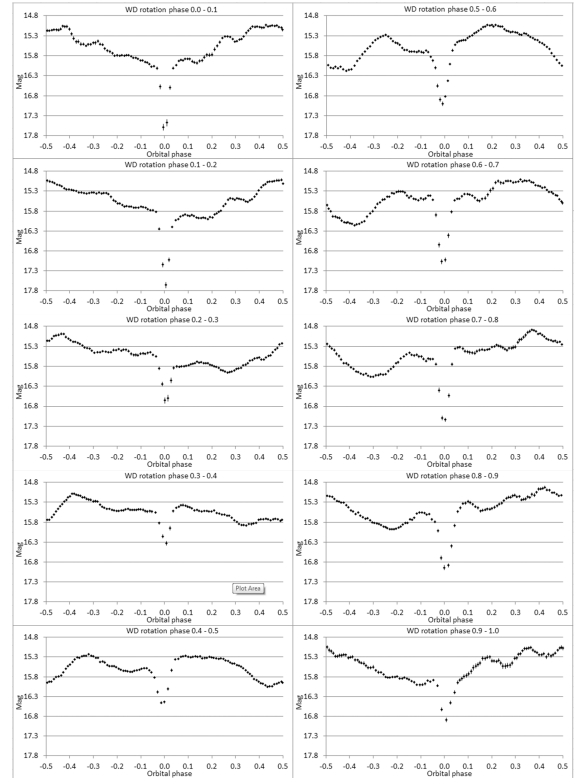


Figure 7. Averaged orbital-phased light curve in 10 bins of WD rotation phase.

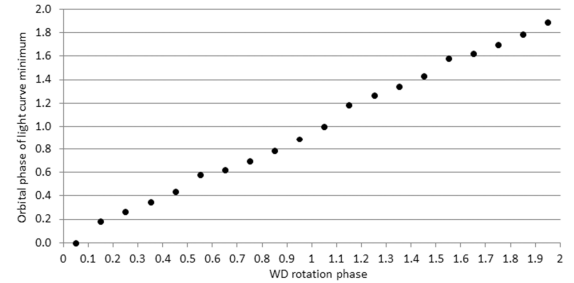


Figure 8. Variation of orbital phase of out-of-eclipse light curve minimum with WD rotation phase.

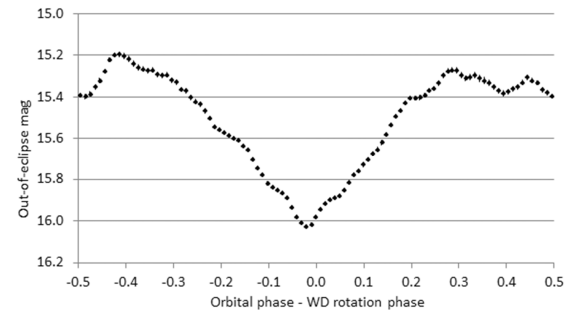


Figure 9. Variation of averaged out-of-eclipse light curve with the difference between orbital and WD rotation phases.

## 7.2 Variation of eclipses with WD rotation phase

Figure 10 shows how the O-C residuals of mid-eclipse times to the linear ephemeris in equation (1) vary with WD rotation phase. Variation with the quadratic ephemeris in equation (2) is virtually identical. Variation of the averaged eclipse profile with WD rotation phase is shown in Figure 11. Eclipses in the visual waveband are generally round-bottomed so either they are grazing or an extended light source is being eclipsed or both. This is in marked contrast to the steep-sided total eclipses observed in the hard X-rays originating in the relatively compact accretion regions close to the magnetic poles of the WD (Mukai et al. 2003). However we note that we find the O-C timing residuals of optical and X-ray eclipses to be fully consistent.

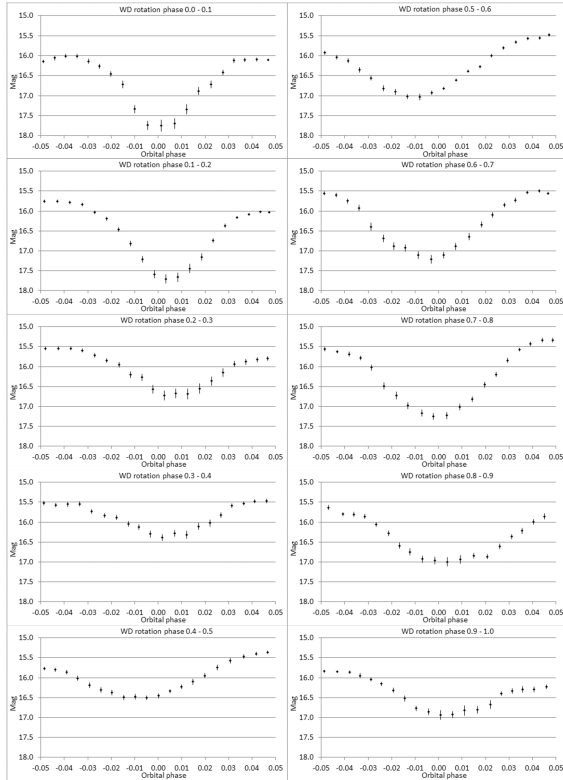


Figure 11. Averaged eclipse profile in 10 bins of WD rotation phase.

As well as variation in mid-eclipse time with WD rotation phase there are also noticeable variations in eclipse depth and width. To examine these variations, we first calculated the mean eclipse depth for each of 10 bins of WD rotation phase by averaging the magnitudes on either side of each eclipse in Figure 7 and subtracting this from the

average magnitude at eclipse minimum calculated from the previous quadratic fits to the eclipses in that bin. We then used a quadratic fit to the lower half of each eclipse in Figure 7 to calculate the mean eclipse width at half-depth for each of 10 bins of WD rotation phase. In Figure 12 we show how eclipse depth, magnitude at eclipse minimum and eclipse width vary with WD rotation phase.

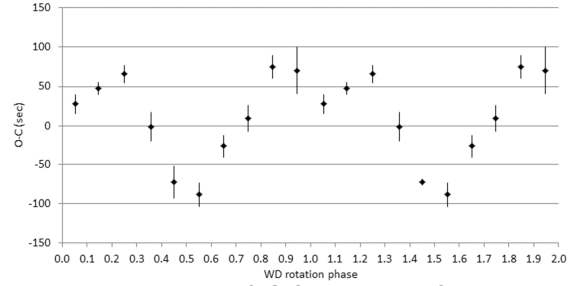


Figure 10. Variation of O-C residuals of mid-eclipse times to the linear eclipse ephemeris in equation (1) with WD rotation phase.

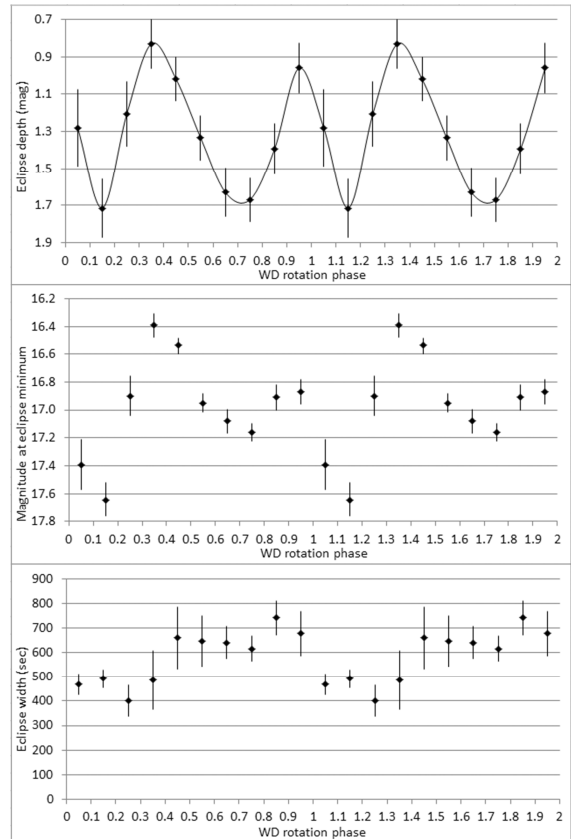
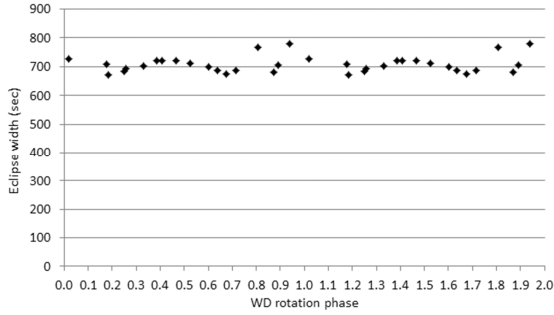


Figure 12. Variation of eclipse depth (upper), magnitude at eclipse minimum (middle) and eclipse width (lower) with WD rotation phase (the curve is purely to help guide the eye).



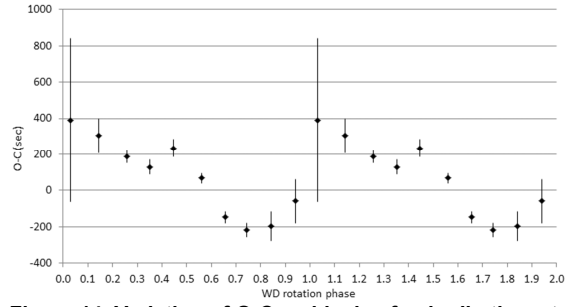
Eclipse depth shows two cycles of variation per WD rotation with maximum amplitude 0.9 magnitudes corresponding to a flux drop of 56%. Eclipse width shows one cycle of variation per WD rotation with a mean width of 463s in the phase range 0.0 to 0.4 and 664s in the range 0.4 to 1.0. This is consistent with the average width of  $632 \pm 35$ s given in Patterson et al. (1995). Eclipse widths measured in X-rays are generally larger than those measured optically. Figure 13 shows durations of 18 X-ray eclipses given in Mukai et al. (2003) for which we have computed the WD rotation phase. There is a degree of correlation in that both optical and X-ray eclipse widths are larger in the phase range 0.8 to 1.0. Note that WD rotation phase 0.0 adopted by Mukai et al. corresponds to phase 0.42 in our analysis.



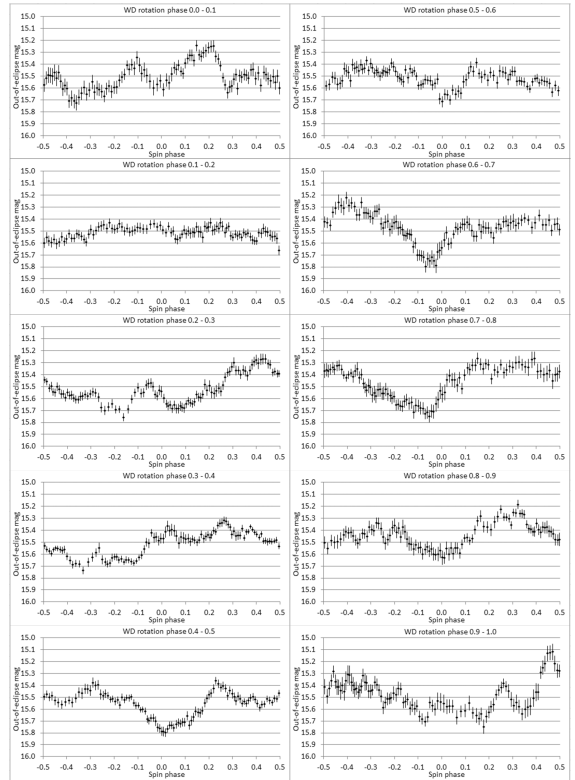
**Figure 13. Variation of X-ray eclipse width from Mukai et al. (2003) with WD rotation phase.**

### 7.3 Variation of WD spin dip with WD rotation phase

Figure 14 shows how the O-C residuals of spin dip times to the best fit ephemeris in equation (3) vary with WD rotation phase. Staubert et al. (2003) considered this variation to be a consequence of both the changing trajectory of the accretion stream in the magnetic field of the rotating WD and movement of its impact point on the WD surface. Figure 15 shows the averaged spin-phased out-of-eclipse light curve in each of 10 bins of WD rotation phase. The most noticeable feature is a persistent minimum around spin phase 0.0 between WD rotation phases 0.4 and 0.9.



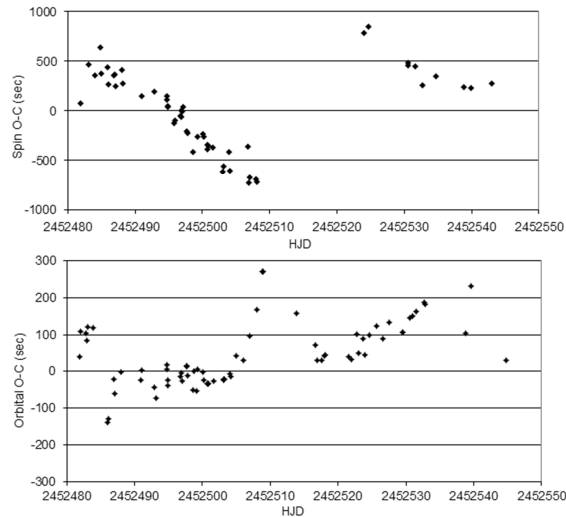
**Figure 14. Variation of O-C residuals of spin dip times to the best fit spin dip ephemeris in equation (3) with WD rotation phase.**



**Figure 15. Averaged spin-phased out-of-eclipse light curve in 10 bins of WD rotation phase.**

## 8. Strange behaviour in 2002

The O-C residuals of spin dip times exhibit a strange discontinuity during 2002 which is shown in Figure 16.



**Figure 16. Spin dip and orbital eclipse O-C residuals in 2002.**

There is also a short-lived increase in the orbital eclipse O-C residuals around the same time. The gap in the spin dip residuals is due to the temporary coincidence of the spin dip with the orbital eclipse which prevents measurement of spin dip timings. These plots include observations from several observers so this is a real effect. On four other occasions during observing seasons in 2007, 2011 (twice) and 2013 there is a coincidence of spin dip and orbital eclipse and thus a gap in the data but in none of these is there a discontinuity in the spin dip O-C residuals like the one in 2002. At the moment we do not have an explanation for this behaviour.

## 9. Summary

Our analysis of 21 years of data on the asynchronous polar V1432 Aql has revealed the following:

- the orbital period of the binary system is 0.140234751d (12116.282s);
- a quadratic ephemeris with  $dP_{orb}/dt = -1.38 \times 10^{-11}$  years/year is slightly favoured over a linear ephemeris;
- the apparent WD spin period has reduced from 0.1406320d (12150.61s) in October 1992 to 0.1405542d (12143.88s) in September 2013;
- the average rate of change of the WD spin period is  $-1.042 \times 10^{-8}$  years/year;

- the WD rotation period has increased from 49.641d to 61.703d over the same interval;
- resynchronisation is expected to occur somewhere between 2097 and 2106;
- there is a dependency of several observable quantities on the WD rotation phase, specifically:
  - orbital-phased light curves;
  - mid-eclipse times;
  - eclipse depth;
  - eclipse width;
  - spin-phased out-of-eclipse light curves;
  - spin dip times.

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Orbital cycle	Mid-eclipse time (HJD)	Assigned uncertainty (d)
12424	2450941.97131	0.00170
12445	2450944.91238	0.00170
12452	2450945.89407	0.00170
12459	2450946.87684	0.00170
12495	2450951.92467	0.00170
12823	2450997.92149	0.00170
13370	2451074.63250	0.00170
13413	2451080.66227	0.00170
15119	2451319.90191	0.00170
15133	2451321.86496	0.00170
15197	2451330.84118	0.00170
17736	2451686.89599	0.00170
18640	2451813.66965	0.00120
18649	2451814.93139	0.00060
18784	2451833.86321	0.00120
18806	2451836.94880	0.00060
23405	2452481.88785	0.00120
23406	2452482.02892	0.00120
23412	2452482.87027	0.00120
23413	2452483.01027	0.00120
23414	2452483.15093	0.00120
23419	2452483.85208	0.00120
23434	2452485.95261	0.00120
23435	2452486.09296	0.00120
23441	2452486.93560	0.00120
23442	2452487.07539	0.00120
23449	2452488.05770	0.00120
23469	2452490.86214	0.00120
23470	2452491.00269	0.00120
23484	2452492.96545	0.00120
23485	2452493.10534	0.00120
23497	2452494.78907	0.00050
23497	2452494.78920	0.00120
23498	2452494.92880	0.00120
23498	2452494.92895	0.00050
23511	2452496.75214	0.00050
23512	2452496.89248	0.00050
23513	2452497.03245	0.00120
23518	2452497.73407	0.00120
23518	2452497.73409	0.00050
23519	2452497.87402	0.00050
23524	2452498.57475	0.00110
23525	2452498.71557	0.00110
23528	2452499.13566	0.00060
23529	2452499.27658	0.00060
23535	2452500.11789	0.00060
23536	2452500.25787	0.00060
23539	2452500.67848	0.00050
23539	2452500.67848	0.00120
23540	2452500.81869	0.00120
23540	2452500.81871	0.00050
23546	2452501.66021	0.00110
23556	2452503.06257	0.00120
23556	2452503.06261	0.00060
23557	2452503.20285	0.00060
23563	2452504.04442	0.00060
23564	2452504.18457	0.00060
23570	2452505.02661	0.00120
23577	2452506.00811	0.00120
23584	2452506.99055	0.00120
23591	2452507.97303	0.00120
23597	2452508.81561	0.00120
23598	2452508.95584	0.00120
23633	2452513.86275	0.00120

Orbital cycle	Mid-eclipse time (HJD)	Assigned uncertainty (d)
23653	2452516.66647	0.00080
23655	2452516.94642	0.00120
23660	2452517.64760	0.00080
23662	2452517.92821	0.00120
23663	2452518.06848	0.00120
23688	2452521.57430	0.00080
23691	2452521.99492	0.00120
23696	2452522.69689	0.00120
23698	2452522.97676	0.00120
23703	2452523.67840	0.00080
23705	2452523.95835	0.00120
23710	2452524.66017	0.00080
23717	2452525.64210	0.00080
23724	2452526.62333	0.00110
23731	2452527.60549	0.00110
23745	2452529.56847	0.00110
23745	2452529.56847	0.00080
23752	2452530.55056	0.00110
23755	2452530.97133	0.00120
23759	2452531.53239	0.00110
23767	2452532.65455	0.00120
23768	2452532.79474	0.00120
23811	2452538.82392	0.00120
23817	2452539.66681	0.00120
23854	2452544.85315	0.00060
36376	2454300.87311	0.00110
36472	2454314.33505	0.00120
36479	2454315.31708	0.00120
36480	2454315.45684	0.00120
36486	2454316.29815	0.00120
36487	2454316.43881	0.00120
36493	2454317.28026	0.00120
36507	2454319.24369	0.00120
36508	2454319.38380	0.00120
36514	2454320.22538	0.00120
36521	2454321.20711	0.00120
36522	2454321.34760	0.00120
36529	2454322.32891	0.00120
36530	2454322.46942	0.00120
36536	2454323.31105	0.00120
36543	2454324.29277	0.00120
36544	2454324.43305	0.00120
36550	2454325.27433	0.00120
36551	2454325.41490	0.00120
36557	2454326.25544	0.00120
36558	2454326.39607	0.00120
36561	2454326.81610	0.00110
36564	2454327.23696	0.00120
36586	2454330.32243	0.00120
36587	2454330.46293	0.00120
36593	2454331.30414	0.00120
36594	2454331.44441	0.00120
36607	2454333.26534	0.00120
36611	2454333.82602	0.00110
36621	2454335.23023	0.00120
36629	2454336.35290	0.00120
36664	2454341.25801	0.00120
36665	2454341.39846	0.00120
36672	2454342.38032	0.00120
36679	2454343.36197	0.00120
36686	2454344.34412	0.00120
36700	2454346.30739	0.00120
36701	2454346.44747	0.00120
36736	2454351.35576	0.00120

Orbital cycle	Mid-eclipse time (HJD)	Assigned uncertainty (d)
36774	2454356.68560	0.00110
36781	2454357.66765	0.00110
36793	2454359.35073	0.00120
36795	2454359.63099	0.00110
36802	2454360.61279	0.00110
36817	2454362.71860	0.00110
36828	2454364.25867	0.00120
46631	2455738.98035	0.00120
46638	2455739.96194	0.00120
46653	2455742.06235	0.00120
46706	2455749.49564	0.00120
46706	2455749.49577	0.00140
46707	2455749.63560	0.00120
46719	2455751.31885	0.00120
46720	2455751.45882	0.00120
46720	2455751.45966	0.00140
46721	2455751.59893	0.00120
46727	2455752.44007	0.00120
46728	2455752.58004	0.00120
46728	2455752.58025	0.00140
46735	2455753.56188	0.00120
46748	2455755.38592	0.00120
46749	2455755.52527	0.00120
46756	2455756.50686	0.00120
46836	2455767.72743	0.00110
46857	2455770.67299	0.00110
46858	2455770.81266	0.00110
46865	2455771.79521	0.00110
46871	2455772.63696	0.00110
46872	2455772.77764	0.00110
46886	2455774.74127	0.00110
46936	2455781.75133	0.00110
46940	2455782.31240	0.00120
47228	2455822.69791	0.00170
47341	2455838.54598	0.00100
47461	2455855.37520	0.00100
47484	2455858.60066	0.00200
47489	2455859.30139	0.00100
47527	2455864.62802	0.00200
47534	2455865.60823	0.00200
47541	2455866.59119	0.00200
49127	2456089.00534	0.00120
49128	2456089.14559	0.00120
49134	2456089.98689	0.00120
49136	2456090.26786	0.00120
49143	2456091.24949	0.00120
49149	2456092.09085	0.00120
49150	2456092.23124	0.00120
49227	2456103.02993	0.00120
49234	2456104.01062	0.00120
49247	2456105.83402	0.00330
49249	2456106.11507	0.00120
49264	2456108.21419	0.00120
49276	2456109.89809	0.00330
49277	2456110.03923	0.00120
49278	2456110.17889	0.00120
49311	2456114.80858	0.00330
49325	2456116.77011	0.00330
49368	2456122.80118	0.00330
49397	2456126.86745	0.00330
49425	2456130.79598	0.00330
49430	2456131.49602	0.00110
49437	2456132.47844	0.00110
49444	2456133.46078	0.00110

Orbital cycle	Mid-eclipse time (HJD)	Assigned uncertainty (d)
49451	2456134.44161	0.00110
49468	2456136.82462	0.00330
49480	2456138.51058	0.00110
49615	2456157.44091	0.00110
49629	2456159.40428	0.00110
49636	2456160.38579	0.00110
49757	2456177.35238	0.00110
49921	2456200.35192	0.00110
51741	2456455.57996	0.00100
51755	2456457.54278	0.00100
51764	2456458.80419	0.00110
51791	2456462.59118	0.00100
51798	2456463.57281	0.00100
51876	2456474.51163	0.00100
51877	2456474.65204	0.00100
51940	2456483.48445	0.00100
51941	2456483.62459	0.00100
51948	2456484.60660	0.00100
51949	2456484.74633	0.00100
51950	2456484.88675	0.00100
51956	2456485.72894	0.00100
51963	2456486.70965	0.00100
51964	2456486.85030	0.00100
51970	2456487.69159	0.00100
51978	2456488.81329	0.00110
52134	2456510.69206	0.00110
52261	2456528.50230	0.00100
52269	2456529.62250	0.00110
52276	2456530.60530	0.00110
52277	2456530.74518	0.00110
52339	2456539.43858	0.00100
52346	2456540.42017	0.00100
52383	2456545.61210	0.00220
52412	2456549.67795	0.00220
52583	2456573.65734	0.00110
52746	2456596.51577	0.00110

**Table 2. Orbital cycle numbers, mid-eclipse times and assigned uncertainties.**

Spin cycle	Spin dip time (HJD)	Assigned uncertainty (d)
0	2450944.98512	0.00520
7	2450945.96734	0.00520
638	2451034.69852	0.00520
922	2451074.61808	0.00520
2652	2451317.88399	0.00520
2666	2451319.84949	0.00520
2744	2451330.81220	0.00520
6178	2451813.66523	0.00260
6187	2451814.92921	0.00260
6343	2451836.87035	0.00260
6357	2451838.83669	0.00260
6272	2451826.88156	0.00260
6329	2451834.89853	0.00260
6336	2451835.88560	0.00260
6350	2451837.85438	0.00260
6400	2451844.87992	0.00260
10939	2452483.06526	0.00470
10946	2452484.04821	0.00470
10952	2452484.89503	0.00470
10953	2452485.03256	0.00470
10959	2452485.87689	0.00470
10960	2452486.01551	0.00470
10966	2452486.86014	0.00470
10967	2452487.00087	0.00470
10968	2452487.14011	0.00470
10974	2452487.98550	0.00470
10975	2452488.12455	0.00470
10995	2452490.93498	0.00470
11009	2452492.90387	0.00470
11022	2452494.73064	0.00470
11022	2452494.73102	0.00470
11023	2452494.87045	0.00470
11023	2452494.87052	0.00470
11030	2452495.85265	0.00470
11031	2452495.99358	0.00470
11036	2452496.69711	0.00470
11037	2452496.83754	0.00470
11037	2452496.83829	0.00470
11038	2452496.97882	0.00470
11039	2452497.11997	0.00470
11043	2452497.67944	0.00470
11044	2452497.81985	0.00470
11050	2452498.66128	0.00470
11054	2452499.22543	0.00470
11060	2452500.06934	0.00470
11061	2452500.20953	0.00470
11065	2452500.77043	0.00470
11065	2452500.77097	0.00470
11066	2452500.91146	0.00470
11071	2452501.61430	0.00470
11081	2452503.01741	0.00470
11082	2452503.15864	0.00470
11088	2452504.00387	0.00470
11089	2452504.14228	0.00470
11108	2452506.81640	0.00470
11109	2452506.95272	0.00470
11110	2452507.09395	0.00470
11116	2452507.93727	0.00470
11117	2452508.07753	0.00470
11230	2452523.98226	0.00470
11235	2452524.68604	0.00470
11277	2452530.58654	0.00470
11277	2452530.58682	0.00470
11284	2452531.57057	0.00470

Spin cycle	Spin dip time (HJD)	Assigned uncertainty (d)
11292	2452532.69315	0.00470
11306	2452534.66247	0.00470
11336	2452538.87909	0.00470
11343	2452539.86317	0.00470
11365	2452542.95679	0.00470
23838	2454296.47818	0.00130
23845	2454297.46412	0.00130
23869	2454300.83790	0.00130
24081	2454330.64265	0.00130
24103	2454333.73359	0.00130
24288	2454359.74021	0.00130
24266	2454356.64910	0.00130
24029	2454323.32964	0.00130
24043	2454325.30222	0.00130
24044	2454325.43949	0.00130
24051	2454326.42521	0.00130
24057	2454327.26776	0.00130
24079	2454330.36156	0.00130
24085	2454331.20542	0.00130
24086	2454331.34656	0.00130
24087	2454331.48360	0.00130
24093	2454332.33067	0.00130
24094	2454332.46977	0.00130
24101	2454333.45237	0.00130
24114	2454335.27992	0.00130
24115	2454335.42167	0.00130
24121	2454336.26490	0.00130
24122	2454336.40461	0.00130
24157	2454341.32395	0.00130
24164	2454342.30913	0.00130
24171	2454343.29325	0.00130
24172	2454343.43382	0.00130
24179	2454344.41706	0.00130
24186	2454345.40135	0.00130
24193	2454346.38649	0.00130
24228	2454351.30712	0.00130
24242	2454353.27341	0.00130
24278	2454358.33464	0.00130
24285	2454359.31780	0.00130
24313	2454363.25394	0.00130
26328	2454646.51335	0.00480
26498	2454670.41927	0.00480
34100	2455739.01848	0.00190
34107	2455740.00072	0.00190
34121	2455741.97212	0.00190
34122	2455742.10940	0.00190
34105	2455739.71937	0.00190
34362	2455775.84363	0.00190
34304	2455767.69155	0.00190
34326	2455770.78386	0.00190
34468	2455790.74504	0.00190
34174	2455749.42039	0.00190
34175	2455749.55986	0.00190
34189	2455751.52421	0.00190
34196	2455752.51110	0.00190
34203	2455753.49393	0.00190
34217	2455755.46037	0.00190
34224	2455756.44502	0.00190
34225	2455756.58561	0.00190
34196	2455752.51121	0.00190
34695	2455822.65156	0.00190
34724	2455826.72833	0.00190
34943	2455857.51562	0.00190
34956	2455859.34154	0.00190

Spin cycle	Spin dip time (HJD)	Assigned uncertainty (d)
34922	2455854.56295	0.00190
34951	2455858.63794	0.00190
34979	2455862.57258	0.00190
36596	2456089.86108	0.00240
36703	2456104.90021	0.00240
36710	2456105.88315	0.00240
36823	2456121.76830	0.00240
36831	2456122.89056	0.00240
36845	2456124.85722	0.00240
36598	2456090.14392	0.00240
36605	2456091.12676	0.00240
36606	2456091.26657	0.00240
36611	2456091.97001	0.00240
36668	2456099.98093	0.00240
36711	2456106.02548	0.00240
36712	2456106.16508	0.00240
36725	2456107.99490	0.00240
36732	2456108.98033	0.00240
36733	2456109.11594	0.00240
36739	2456109.95994	0.00240
36740	2456110.09998	0.00240
36899	2456132.44800	0.00240
36913	2456134.41534	0.00240
39397	2456483.55915	0.00180
39404	2456484.54363	0.00180
39716	2456528.39909	0.00180
39333	2456474.56417	0.00180
39362	2456478.63799	0.00180
39434	2456488.76145	0.00180
39221	2456458.82154	0.00180
39399	2456483.84141	0.00180
39442	2456489.88278	0.00180
39718	2456528.68143	0.00180
39725	2456529.66300	0.00180
39732	2456530.64621	0.00180
39405	2456484.68479	0.00180
39406	2456484.82454	0.00180
39412	2456485.66880	0.00180
39419	2456486.65396	0.00180
39420	2456486.79372	0.00180
39427	2456487.77532	0.00180
39696	2456525.58733	0.00180
39839	2456545.68871	0.00180
39846	2456546.67362	0.00180

**Table 3. WD spin cycle numbers, spin dip times and assigned uncertainties.**